Retrieval of H₂O abundance in Titan’s stratosphere: a (re)analysis of CIRS/Cassini and PACS/Herschel observations

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Abstract

Since its first measurement 20 years ago by the Infrared Space Observatory (ISO), the water (H₂O) mole fraction in Titan’s stratosphere remains uncertain due to large differences between the determinations from available measurements. More particularly, the recent measurements made from the Herschel observatory (PACS and HIFI) estimated the H₂O mole fraction to be 0.023 ppb at 12.1 mbar. A mixing ratio of 0.14 ppb at 10.7 mbar was, however, retrieved from nadir spatially-resolved observations of Cassini/CIRS. At the same pressure level (10.7 mbar), this makes a difference of a factor of 5.5 between PACS and CIRS measurements, and this has notably prevented current models from fully constraining the oxygen flux flowing into Titan’s atmosphere. In this work, we try to understand the differences between...
the H$_2$O mole fractions estimated from Herschel/PACS and Cassini/CIRS observations. The strategy for this is to 1) analyse recent disc-averaged observations of CIRS to investigate if the observation geometry could explain the previous observed differences, and 2) (re)analyse the three types of observation with the same retrieval scheme to assess if previous differences in retrieval codes/methodology could be responsible for the previous discrepancies. With this analysis, we show that using the same retrieval method better reconcile the previous measurements of these instruments. However, the addition of the disc-averaged CIRS observations, instead of confirming the consistency between the different datasets, reveals discrepancies between one of the CIRS disc-averaged set of observations and PACS measurements. This raises new questions regarding the possibility of latitudinal variations of H$_2$O, which could be triggered by seasonal changes of the meridional circulation. As it has already been shown for nitriles and hydrocarbons, this circulation could potentially impact the latitudinal distribution of H$_2$O through the subsidence or upwelling of air rich in H$_2$O. The possible influence of spatial/time variations of the OH/H$_2$O input flux in Titan’s atmosphere is also discussed. The analysis of more observations will be needed in future work to address the questions arising from this work and to improve the understanding of the sources of H$_2$O in Titan’s atmosphere.

**Keywords:** Titan, atmosphere, Atmosphere, composition, Spectroscopy, Infrared observations, Satellite, atmosphere

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1. *Herschel is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.*
1. Introduction

Since the discovery of carbon dioxide ($\text{CO}_2$) more than 30 years ago (Samuelson et al., 1983), the presence of oxygen compounds has been firmly demonstrated in Titan’s atmosphere. Among these, water vapour ($\text{H}_2\text{O}$) is deposited in the high atmosphere by external sources. It was measured for the first time in Titan’s stratosphere around 20 years ago by the Infrared Space Observatory (ISO), using the Short Wavelength Spectrometer (SWS) (Coustenis et al., 1998). The two lines recorded at 227.8 and 254 cm$^{-1}$ by this instrument were effectively modelled using a $\text{H}_2\text{O}$ profile of 0.4 ppb above some cut-off altitude and were mainly sensitive to $\text{H}_2\text{O}$ in the pressure range $1–1 \times 10^{-5}$ mbar. After this first detection, de Kok et al. (2007a) made a first attempt to measure stratospheric $\text{H}_2\text{O}$ with the Composite Infrared Spectrometer (CIRS) on board the NASA Cassini spacecraft. Due to a poor signal-to-noise ratio (SNR), this first analysis of CIRS spectra was unsuccessful, and only an upper limit of 0.9 ppb was determined, which was consistent with ISO observations. These results were, however, later challenged by measurements from the Photodetector Array Camera and Spectrometer (PACS) and the Heterodyne Instrument for the Far-Infrared (HIFI) onboard the Herschel observatory (Moreno et al., 2012). Using both instruments, Moreno et al. (2012) constrained the vertical profile of $\text{H}_2\text{O}$, showing that it increases with height, which is consistent with a high altitude source and low-level sink (condensation and photolysis). They retrieved a $\text{H}_2\text{O}$ volume mixing ratio (VMR) of 0.023 ppb at 12.1 mbar, which is 20-times lower than the VMR (assumed constant with altitude above the saturation level) determined from ISO. Furthermore, Moreno et al. (2012) reanalysed SWS/ISO spectra and
showed that the H$_2$O abundance inferred from those observations needed a downward revision to 0.06 ppb (uniform profile). Almost at the same time, Cottini et al. (2012) performed a second attempt to measure H$_2$O from nadir and limb CIRS observations. By averaging a large number of nadir spectra recorded over 4 years, they determined a H$_2$O VMR of 0.14 ppb at 10.7 mbar. From limb observations, they retrieved 0.13 ppb at 115 km (around 6 mbar) and 0.45 ppb at 230 km (around 0.35 mbar), which confirmed that the H$_2$O mole fraction increases with height. The retrieved abundance, however, differed by about a factor of 4 from the results of Moreno et al. (2012).

These uncertainties on the H$_2$O mole fraction have prevented photochemical models from constraining the flux of OH/H$_2$O into Titan’s atmosphere. As a result, it is currently not determined whether the main source of H$_2$O comes from micrometeorite ablation (English et al., 1996) or from local sources, such as the cryovolcanic activity on Enceladus (e.g. Hansen et al., 2006; Waite et al., 2006; Hansen et al., 2008, 2011). Using HIFI onboard Herschel, Hartogh et al. (2011) detected the Enceladus H$_2$O torus and proposed that the latter could be the main source of H$_2$O in Saturn’s atmosphere. Based on the OH/H$_2$O input flux required by the model of Hörst et al. (2008) to reproduce the H$_2$O abundance retrieved from ISO observations, Hartogh et al. (2011) concluded that the Enceladus influx rates at Titan’s atmosphere are too small to explain the observed H$_2$O abundance. However, with the retrieval of H$_2$O from Herschel, Moreno et al. (2012) estimated new input fluxes and showed that Enceladus surface activity is a viable source of H$_2$O in Titan’s atmosphere. This was confirmed by the recent model of Dobrijevic et al. (2014), coupling oxygen, nitrogen and hydrocarbon chemistry,
which was able with an Enceladus source to reproduce the H\textsubscript{2}O abundance retrieved by either Moreno et al. (2012) or Cottini et al. (2012). Note that if the H\textsubscript{2}O abundance retrieved by Cottini et al. (2012) is taken as the reference, Dobrijevic et al. (2014) overestimated the thermospheric H\textsubscript{2}O mole fraction when compared with the upper limit estimated by Cui et al. (2009) from the Cassini/INMS (Ion Neutral Mass Spectrometer) data. The OH/H\textsubscript{2}O input flux determined is, however, very different depending on the reference chosen. More particularly, if the flux is determined using the results of Moreno et al. (2012), both Moreno et al. (2012) and Dobrijevic et al. (2014) could not reproduce H\textsubscript{2}O and CO\textsubscript{2} abundances at the same time. They respectively obtained CO\textsubscript{2} mole fractions 10 and 4 times lower than the observed abundance (de Kok et al., 2007a). Based on the large difference in the atmospheric lifetimes of CO\textsubscript{2} (several hundred years) and H\textsubscript{2}O (around 10 years), Moreno et al. (2012) invoked a variable OH/H\textsubscript{2}O input flux over long timescales to explain this difference. This scenario has been explored by the recent model of Lara et al. (2014). They showed, however, that when results of Moreno et al. (2012) are considered, a time-dependent input flux does not solve alone the H\textsubscript{2}O/CO\textsubscript{2} problem and the addition of another loss term for H\textsubscript{2}O, such as a loss to the haze, is required. Understanding the discrepancies between the H\textsubscript{2}O mole fractions measured from Herschel and CIRS is therefore crucial to better constrain the existing photochemical models, and to improve the knowledge of the source and chemistry of oxygen compounds in Titan’s atmosphere.

In this work, we aim to investigate and understand these previous reported differences. For this, we analyse recent disc-averaged observations
of Titan (2013-2015) made with the CIRS/Cassini instrument and retrieve new H$_2$O mole fractions. The use of disc-averaged observations allows an easier comparison with Herschel measurements, which share the same viewing geometry. In addition to a difference in observation geometry, Moreno et al. (2012) and Cottini et al. (2012) also applied distinct retrieval methods and used different a priori information and radiative transfer modelling codes to retrieve H$_2$O. The goal of this work is also to (re)analyse the CIRS and Herschel datasets with the same retrieval scheme. This will allow us to distinguish if the observed discrepancies are due to modelling/retrieval differences or due to other possible sources. In the case of Herschel, we focus only on PACS observations, which have a very similar vertical sensitivity to H$_2$O as the CIRS measurements. Among the spectra analysed by Cottini et al. (2012), we only consider the nadir spectrum for the same reasons. These observations are presented in more details in section 2. Sections 3 and 4 describe, respectively, the forward modelling settings and the retrieval strategy developed to analyse all observations. The results are then discussed in section 5.

2. Instruments and observations

2.1. CIRS

The CIRS instrument is a Fourier transform spectrometer. It is composed of three separate focal planes (FP1, FP3 and FP4) that measure together spectra in the far- and mid-infrared (IR) (full range: 10–1400 cm$^{-1}$) with an adjustable apodised resolution between 0.5 and 15.5 cm$^{-1}$. The FP1 covers the far-IR spectral range (10–600 cm$^{-1}$). It has a circular field-of-view
(FOV) of 3.9 mrad diameter, whose half of the integrated response is located within 2.5 mrad. The FP3 and FP4 focal planes consist of collinear arrays of 10 pixels each that record spectra in the mid-IR range (respectively 600–1100 cm$^{-1}$ and 1100–1400 cm$^{-1}$). The size of each FOV is 0.27 $\times$ 0.27 mrad.

More details about the instrument can be found elsewhere (e.g. Flasar et al., 2004; Jennings et al., 2017).

In this work, four sets of FP1 observations recorded between February 2013 and July 2015 at a spectral resolution of 0.5 cm$^{-1}$ are analysed. They were acquired during four different observations at more than 1.5 million kilometres from Titan (Titan Explorations at Apoapse, TEA), and each set include around 800 disc-averaged spectra. The first observation (rev 182) lasted 20 hours whereas the three next ones (revs 202, 206, 209) lasted around 12 hours. The sub-spacecraft latitudes of these observations are -41.43$^\circ$, 50.73$^\circ$, 48.58$^\circ$ and -0.02$^\circ$ respectively for sets 182, 202, 206 and 219. Details of these four sets are given in Table 1.

Because H$_2$O lines are very weak in the FP1 range, an averaged spectrum has been calculated for each set of observations to increase the SNR. They are shown in Figure 1 for the 125–260 cm$^{-1}$ spectral range. The associated noise has been evaluated using the same method as Teanby et al. (2006). Briefly, two types of standard error are calculated for each set: 1) the standard deviation on the averaged spectrum, and 2) the standard error related to the noise equivalent spectral radiance (NESR) measured from deep space spectra. The maximum of these standard errors is then taken as the measurement error on the averaged spectrum. This measurement error is similar for the four sets, with the minimum found for the 182 (around 1 nW cm$^{-2}$ sr$^{-1}$ cm
for the spectral range shown in Figure 1) and the maximum for the 206 and 219 sets (around 1.4 nW cm\(^{-2}\) sr\(^{-1}\) cm). Noise spikes are however located in different ranges: 140–150 cm\(^{-1}\) for 182, 152–164 cm\(^{-1}\) for 202, 212–227 cm\(^{-1}\) for 206, and 240–265 cm\(^{-1}\) for 219. A spike at 191.25 cm\(^{-1}\) is common to all sets.

From Figure 1, we can see that H\(_2\)O lines are at most barely detected. This is not surprising given the small number of spectra averaged together. For comparison, Cottini et al. (2012) averaged thousands of nadir spectra to achieve a sufficient SNR. Among the four sets analysed here, the SNR seems to vary, with only set 182 showing clear H\(_2\)O lines. For this set, the SNR of the H\(_2\)O lines is estimated to be around 2 for the lines located above 200 cm\(^{-1}\) and around 3 for the ones below. The difference observed with the other sets is unlikely to be attributable to measurement error differences, the latter being small. The possibility of latitudinal variations will be discussed in section 5.2.4. Note that in section 5.2.1, we will assess if H\(_2\)O lines are present in the sets 202, 206 and 219. Note also that the retrieval scheme, which is described in the following sections, has been developed using set 182 only.

Finally, the nadir spectrum studied previously by Cottini et al. (2012) is reanalysed. This spectrum is an average of around 7000 spectra acquired between December 2004 and December 2008 in a latitudinal bin of 0–30°N (see Figure 1). More details are given in Cottini et al. (2012). This averaged spectrum will be called “CIRS NA” hereafter.
2.2. PACS/Herschel

PACS is one of the three science instruments onboard the Herschel observatory. It is composed of an imaging photometer, covering the spectral range 60–210 μm (47.6–181.8 cm$^{-1}$) over a FOV of 1.75′ × 3.5′, and of a grating spectrometer, providing observations in the range 55–210 μm at a spectral resolving power between 1000 and 4000 and over a FOV of 47″ × 47″, resolved in 5 × 5 pixels. More details about the instruments can be found in Poglitsch et al. (2010).

In this work, we reanalyse the spectra recorded by the PACS spectrometer that were previously studied by Moreno et al. (2012). These consist of disc-averaged observations of three H$_2$O lines at 66.43 μm (150.5 cm$^{-1}$), 75.38 μm (132.7 cm$^{-1}$) and 108.07 μm (92.5 cm$^{-1}$), expressed as line/continuum ratios. They were recorded on 22nd June 2010 when the apparent distance between Titan and Saturn was close to the maximum elongation (see Moreno et al. (2012) for more details about the observations and their reduction). In addition to H$_2$O lines, observations of CH$_4$ lines acquired by PACS at around the same time are also analysed to retrieve temperature information. These were recorded in the range 102–146 μm using the chopped-nodded PACS range spectroscopy mode with a small chopper throw (see PACS observers’ manual 2013, http://herschel.esac.esa.int/Docs/PACS/pdf/pacs_om.pdf). In this spectral range, two CH$_4$ lines have been used in this work, at 106.43 μm (94 cm$^{-1}$) and 119.63 μm (83.6 cm$^{-1}$). These CH$_4$ spectra have been reduced from Level 0 to Level 2 within HIPE 15.0.0 (Herschel Interactive Processing Environment), using the Telescope Normalization pipeline (see PACS Data Reduction Guide 2017 for details, http://herschel.esac.esa.int/twiki/pub/Pu
bic/PacsCalibrationWeb/pacs_spec15.pdf), and the final spectrum has been rebinned close to the Nyquist sampling. Note that spectra (CH\textsubscript{4} lines only) reduced in such a way have units of Janskys (Jy) and have been converted to spectral radiance units of W cm\textsuperscript{-2} sr\textsuperscript{-1} cm. Table 2 summarizes the PACS observations analysed in this work, which are shown in Figure 1.

3. Forward model

3.1. Reference atmosphere and spectroscopic data

A reference atmosphere has been defined for the analysis of all the observations considered. It is composed of 99 levels distributed between 1457 and 1.01 \times 10\textsuperscript{-5} mbar. The far-infrared spectrum of Titan’s atmosphere, and more especially the continuum level, is affected by collision-induced absorption (CIA) between pairs of N\textsubscript{2}, CH\textsubscript{4} and H\textsubscript{2} molecules. The CIA coefficients have been calculated using the work of Borysow and Frommhold (1986a,b,c, 1987), Borysow (1991) and Borysow and Tang (1993). The reference vertical profiles of the 3 gases included are shown in Figure 2.a. The profile of CH\textsubscript{4} has been defined according to Niemann et al. (2010) with a volume mixing ratio of around 5.6\% close to the surface and of 1.48\% in the stratosphere. For H\textsubscript{2}, the mixing ratio has been set at 0.101\% at all pressure levels (Niemann et al., 2010). In addition to CIA, the continuum shape is also impacted by aerosol absorption. The assumed haze vertical profile is shown in Figure 2.b and is a simplified version built from previous results (de Kok et al., 2007b, 2010b; Tomasko et al., 2008). The haze relative absorption cross-sections adopted were previously used by Teanby et al. (2013) and derived from the volume absorption coefficients of Anderson and Samuelson (2011). Finally, emission
lines of different gases are present in the far-infrared range and are mainly those of CO, HCN, CH₄, H₂O, C₄H₂ and C₂N₂. The spectral ranges used in the different steps of our retrieval scheme have, however, been selected to only include either CH₄ and its isotopologues (CH₃D and ¹³CH₄) or H₂O lines. The vertical profiles of CH₃D and ¹³CH₄ we have considered are shown in Figure 2.a and have been calculated using the isotopic ratios determined by Niemann et al. (2010). The choice of the temperature and the H₂O a priori profiles will be discussed in the following sections. Note finally that line data were taken from the HITRAN 2004 database (Rothman et al., 2005) and CH₄ line intensities have been revised according to Wishnow et al. (2007). For CIRS, these have been used to calculate k-tables (Lacis and Oinas, 1991) of gaseous opacities (see Teanby et al. (2013) for details) and line-by-line calculations have been performed for the H₂O retrieval. For PACS, given its high spectral resolution, line-by-line calculations have been preferred to the correlated-k method. To decrease the computation time, pre-calculated tables of monochromatic absorption coefficients (hereafter lbl-tables) have been built using the same pressure/temperature grids as for calculating the k-tables.

3.2. Disc-averaged radiance

To analyse both CIRS FP1 (except the nadir spectrum) and PACS observations, it has been necessary to accurately model the disc-averaged radiance. More especially, in the case of CIRS, the inhomogeneous response of the FP1 detector has to be taken into account. It has been shown that the disc-averaged spectrum can be calculated using a weighted sum of spectra simulated for discrete FOV located at increasing offsets from Titan’s centre.
(Teanby and Irwin, 2007; Teanby et al., 2013). In this work, a total of 20 synthetic spectra distributed from the centre of Titan to the top of the atmosphere has been used to calculate the expected disc-averaged radiance of Titan. The weights applied to the sum are calculated as explained in Teanby et al. (2013). In the case of CIRS observations, these weights are multiplied by the FP1 spatial sensitivity response determined while Cassini was en route to Saturn (Flasar et al., 2004).

4. Analysis strategy

The retrieval of H$_2$O abundance has been performed using the NEMESIS software (Irwin et al., 2008). This code includes a forward model which can be solved either using a line-by-line model or the correlated-k method (Lacis and Oinas, 1991) according to the spectral resolution of the measurement. A retrieval method is also implemented and is based on the optimal estimation formalism (Rodgers, 2000). The idea of this inverse method is to determine the most probable atmospheric state that is consistent with both the spectrum and the knowledge of the atmosphere prior to the measurement. This is done by minimising a cost function $\phi$ of the form

$$
\phi = \left( \mathbf{y} - \mathbf{F}(\mathbf{x}) \right)^T \mathbf{S}_\epsilon^{-1} \left( \mathbf{y} - \mathbf{F}(\mathbf{x}) \right) + \left( \mathbf{x} - \mathbf{x}_a \right)^T \mathbf{S}_a^{-1} \left( \mathbf{x} - \mathbf{x}_a \right),
$$

where $\mathbf{y}$ is the measured spectrum, $\mathbf{x}$ is the atmospheric state vector, $\mathbf{x}_a$ is the a priori state vector, $\mathbf{S}_\epsilon$ is the measurement covariance matrix (including measurement errors), $\mathbf{S}_a$ is the a priori covariance matrix (including the uncertainties on the a priori state), and $\mathbf{F}$ is the forward model. The
minimisation is performed using an iterative scheme based on the Levenberg-
Marquardt method (see Irwin et al. (2008) for more details). Note that in
this work, the a priori uncertainties are chosen to avoid unconstrained (ele-
ments of $S_a$ are too large) or over-constrained (elements of $S_a$ are too small)
solutions (Irwin et al., 2008).

To analyse CIRS and PACS observations, we have developed a retrieval
scheme relying on three successive steps: 1) the retrieval of the tropospheric
and stratospheric temperatures, 2) the adjustment of the continuum level in
the vicinity of the targeted H$_2$O lines, and 3) the retrieval of H$_2$O strato-
spheric abundance. Because the stratospheric temperatures and the contin-
um level affect the retrieval of H$_2$O abundance (see below), their associated
uncertainties have been evaluated and then propagated to evaluate the total
error on the retrieved H$_2$O abundance. Each step of the analysis is described
in more details in the following sections.

4.1. Step 1: Temperature retrieval

Because it affects the strength of H$_2$O emission lines along with the H$_2$O
abundance itself, stratospheric temperatures have to be retrieved. As well as
affecting the strength of H$_2$O emission lines, the temperature profile also im-
pacts the vertical sensitivity to H$_2$O. The lower boundary of this sensitivity
is indeed determined by the condensation level of H$_2$O, whose height depends
on the assumed temperature profile. Colder temperatures move this level to-
wards higher altitudes and inversely. For these two reasons, the temperature
profile has to be determined before the H$_2$O retrieval, and more especially in
the pressure range of maximal sensitivity to H$_2$O. From the H$_2$O Jacobians$^2$
shown in Figure 3 (b,d), this range is found to be between 12 and 1 mbar for
both disc-averaged CIRS and PACS observations (Jacobians calculated for
the CIRS NA spectrum are similar to those of CIRS disc-averaged observa-
tions). The temperature over this range of pressure levels is usually retrieved
for CIRS observations using the CH$_4$ $\nu_4$ band of spatially-resolved observa-
tions recorded with the FP4 focal planes (e.g. Achterberg et al., 2008, 2011;
Vinatier et al., 2010). If these observations are used to build a disc-averaged
temperature profile, they have to sample a sufficiently large range of different
latitudes and have to be recorded in a same time interval, close to the time of
the FP1 disc-averaged observations. This was found to be difficult to achieve.
Therefore, we decided to retrieve the whole temperature profile from the far-
IR spectrum only. Pure rotational CH$_4$ lines along with small parts of the
continuum level in their vicinity have been used to retrieve stratospheric and
tropospheric temperatures. Although this spectral range is mainly sensitive
to tropospheric temperatures (Figure 3.a and 3.c), we can see from Figure 3
that the temperature Jacobians calculated for the pure CH$_4$ rotational lines
for both instruments are not null in the 12-1 mbar range, especially close to
the level where the Jacobians of H$_2$O are the largest. CH$_4$ rotational lines
are thus partly sensitive to temperatures in this region and this is what has
been exploited in the retrieval. The vertical sensitivity of the temperature
retrieval will be discussed in section 5.1. In the case of the CIRS NA spec-
trum, the temperature is not disc-averaged but has been retrieved using the

$^2$Derivatives of the calculated top-of-atmosphere radiance with respect to the H$_2$O
abundances at the 99 levels of the reference atmosphere
same method. In the next subsections, the retrieval settings are described and the method used to assess the errors on the retrieved temperature profile is presented.

4.1.1. Retrieval settings

For CIRS FP1 observations, the continuous temperature profile has been retrieved using six CH\textsubscript{4} lines, which have the largest sensitivity to stratospheric temperature (Figure 5), and including the continuum level in their vicinity (Figure 5). In the case of PACS, fewer lines were available and the temperature has been retrieved from the lines located at 106.43 μm (94 cm\textsuperscript{-1}) and 119.63 μm (83.6 cm\textsuperscript{-1}). In both cases, the retrieval range has been chosen to not include any other lines than CH\textsubscript{4}.

Depending on the type of observation, we have considered different \textit{a priori} temperature profiles taking into account the time and the sub-spacecraft latitude at which they were recorded. For CIRS NA and PACS spectra, the \textit{a priori} temperature profile has been set to the temperature profiles respectively assumed by Cottini et al. (2012) and Moreno et al. (2012). Cottini et al. (2012) adopted the temperature profile determined by Anderson and Samuelson (2011) at 15°N. The profile used by Moreno et al. (2012) was a combination of 1) the temperatures measured by the Huygens Atmospheric Structure Instrument (HASI) between 0 km and 140 km (Fulchignoni et al., 2005), 2) the disc-averaged temperatures measured from CIRS (140–500 km) by Vinatier et al. (2010), 3) temperatures retrieved from INMS (1000–1500 km) (De La Haye et al., 2007), and 4) a decreasing temperature from 165 K to 155 K between 500 km and 1000 km. For the disc-averaged CIRS observations, the \textit{a priori} temperature profile has been built using
retrieved profiles from nadir and limb mid-IR CIRS observations (see Figure 6) (Teanby et al., 2016). Briefly, temperature has been retrieved at four pressure levels (5.6, 1 mbar from nadir, 0.1 and 0.01 mbar from limb) and interpolated in time and space to acquire a more global coverage. The temperatures estimated at the sub-spacecraft latitude and at the time of the disc-averaged observations have been interpolated onto a finer pressure grid and have been joined onto the HASI profile (Fulchignoni et al., 2005). The a priori uncertainty has been taken at 1% on the whole profile for all the observations.

The measurement covariance matrix $S_\epsilon$ has been considered diagonal and includes the measurement noise associated with each observations. For CIRS observations, the retrieval has been performed using the correlated-k method. Lbl-tables have been used for PACS spectra. To take into account slight differences between these methods and line-by-line calculations, forward modelling errors have been added to the diagonal elements of the $S_\epsilon$. They have been estimated from differences between spectra simulated with the different methods using the reference atmosphere described in section 2.

4.1.2. Error analysis

When using the optimal estimation method, the assessment of the retrieval uncertainties is usually performed using the total error covariance matrix, $S = (K^T S_\epsilon^{-1} K + S_a^{-1})^{-1}$, whose square roots of the diagonal elements correspond to the standard deviation of each retrieved variable included in the retrieved state vector $x_{\text{ret}}$ (i.e., temperature at each pressure level). The use of these errors implies that the elements of the a priori covariance matrix are very well known and representative of the real uncertainties on the
a priori profile, which is not the case here. The retrieval random errors have therefore been evaluated using another method. Here, 35 different temperature profiles, covering a large range of temperatures, have been built for each observation. Using these profiles and the reference atmosphere defined in section 3.1, synthetic spectra have been generated and random noise has been added (assuming a normal distribution with a standard deviation equal to the noise associated to the observation). The retrieval procedure described in the previous subsection has then been applied to the 35 spectra and the error has been estimated from the differences between the retrieved and real temperature profiles.

The systematic error related to the choice of the CH\textsubscript{4} profile used in the retrieval has not been estimated from the above analysis. As already mentioned, the CH\textsubscript{4} lines are also sensitive to CH\textsubscript{4} abundance and the latter has been fixed to the profile determined by Niemann et al. (2010). The uncertainties on this profile propagate to the retrieved temperature profile and have to be taken into account. To assess the error propagation, we have shifted the CH\textsubscript{4} profile by \(\pm 11\%\) (Niemann et al., 2010) and retrieved the temperature profile again. The systematic errors are taken as the temperature difference between the retrieval with and without applying the shift. It is also worth mentioning that the CH\textsubscript{4} profile determined by Niemann et al. (2010) is measured at one specific location and is not necessarily representative of the observations analysed here. More especially, Lellouch et al. (2014) have observed latitudinal variations of CH\textsubscript{4} stratospheric mole fraction and have estimated that the latter can be as low as 0.01 at some latitudes. For the sake of completeness we have also evaluated the difference in the retrieved
temperature profile if the true stratospheric CH$_4$ mole fraction is equal to 0.01. These different effects will be discussed in section 5.1. Note finally that there is some error on the intensities of the CH$_4$ pure rotational lines that we have considered. More especially, there is a 10% reduction of the CH$_4$ line intensities between Wishnow et al. (2007) and the more recent values reported by Boudon et al. (2010). However, we have evaluated that this difference leads to a maximum uncertainty of 0.06 K on the temperature profile. This uncertainty is insignificant compared with the other sources of error described above, and will be therefore neglected in the calculation of the total uncertainty on the retrieved H$_2$O abundance.

4.2. Step 2: Continuum

4.2.1. CIRS

The second step of the retrieval scheme consists of properly modelling the continuum level, especially in the vicinity of H$_2$O lines. Indeed, an under(over)-estimation of this level leads to an over(under)-estimation of the H$_2$O abundance. This level is mainly affected by tropospheric temperatures, aerosols distribution and absorption, and by the CIA between pairs of N$_2$, CH$_4$ and H$_2$ molecules. The retrieval of the temperature profile described in previous section has partly adjusted this level but we have noticed that differences remain between the simulated and the observed spectra, especially above 120 cm$^{-1}$ where most of the targeted H$_2$O lines are located. To reduce these differences, we have retrieved a continuous profile of aerosols in a first stage. The $a$ priori profile considered is shown in Figure 2.b and the related uncertainties have been set to 20%. The retrieval has been performed using the correlated-k method for the spectral ranges 70-100 cm$^{-1}$ and 130-245
cm$^{-1}$, avoiding emission lines of gases and noise spikes. As for temperature, forward modelling errors have been included in the $S_e$ matrix.

The retrieval of aerosols profile brings a slight improvement to the fit of the continuum level but the latter remains unsatisfactory above 120 cm$^{-1}$. We have assumed that this was due to uncertainties in the spectral dependence of the absorption cross-sections of haze that we have considered. To properly fit the observed continuum, keeping the retrieved aerosol profile, we have smoothly adjusted the spectral variations of the aerosol absorption cross-sections in the range 120–255 cm$^{-1}$. Note that apart from variations of haze properties, another possible source for the observed continuum residual is the deficiencies in the modelled CIA (de Kok et al., 2010a), which have not been considered here. However, both aerosols and CIA contribute to the radiance mostly below the emitting level of H$_2$O and, as long as the continuum level is properly fitted, their values should therefore not impact the H$_2$O retrieval. Example fits to the CIRS spectra are shown in Figure 5.

4.2.2. PACS/Herschel

Because they are expressed as line/continuum ratios, the analysis performed on CIRS spectra does not need to be applied to PACS H$_2$O observations. These ratios have however to be converted into disc-averaged radiances. For this, a continuum level has been simulated using the reference atmosphere and using the temperature profile retrieved after the first step. The ratios have then been rescaled using the modelled continuum.
4.3. Step 3: $\text{H}_2\text{O}$ retrieval

4.3.1. Retrieval settings and a priori choice

The last step of the retrieval scheme consists of determining the stratospheric $\text{H}_2\text{O}$ abundance. For CIRS observations, this is done using 5 $\text{H}_2\text{O}$ lines: 132.75, 170.25, 202.75, 208.5, 254 cm$^{-1}$ (identified in Figure 1 and Figure 5), using line-by-line calculations. In the case of PACS, only three lines are used (described in Table 2) and the retrieval is performed using lbl-tables.

According to the optimal estimation method, the \textit{a priori} $\text{H}_2\text{O}$ profile should represent the best knowledge of the $\text{H}_2\text{O}$ abundance prior to the measurement. The choice of this profile is important as it will affect directly the $\text{H}_2\text{O}$ retrieved abundance. This choice is, however, in this case difficult because the uncertainties on the $\text{H}_2\text{O}$ abundance are large, and their reduction is at the centre of this work. Table 3 and Figure 4 present several profiles that could be used as \textit{a priori}. They are either consistent with previous CIRS observations (Cottini et al., 2012) or with the profile determined from Herschel measurements (Moreno et al., 2012). Because it is impossible to establish which of these is the most representative of the true profile, all of them have been considered as \textit{a priori}. This allows us 1) to assess the effect of the \textit{a priori} profile on the retrieval of $\text{H}_2\text{O}$ abundance, and 2) to determine if one of them is better suited. Note that we have considered only profiles that differ in values and shape. We have therefore not considered the profiles modelled by Dobrijevic et al. (2014) assuming an external source of CO, which are similar to the “IM1”, “IE1” and “IM2” profiles listed in Table 3. The $\text{H}_2\text{O}$ profiles simulated by Lara et al. (2014) for the comet
impact scenario have not been used because this scenario is very unlikely. To
take into account the condensation of H$_2$O in the low atmosphere, H$_2$O has
been set to its saturation vapor pressure\(^3\) below the condensation level for
all the \textit{a priori} profiles. A scaling factor, set at 1 as a first guess, is retrieved
for each of them considering 100\% of \textit{a priori} uncertainties.

\subsection*{4.3.2. Error analysis - CIRS}

The retrieved H$_2$O abundance is associated to two types of error: 1) the
error on the abundance itself, and 2) the uncertainty on the height at which
this abundance is located. The first one has been evaluated by propagating
the uncertainties of different parameters, and this has been done for the
retrieval case using the “IM2” H$_2$O profile as \textit{a priori}. We have first con-
sidered the propagation of the random errors on the retrieved temperature
profile and of the measurement noise. For this, the temperature random
errors are incorporated in the measurement covariance matrix $S_\epsilon$ as follows:

$$
S_\epsilon' = S_\epsilon + K_{Tp} S_{Tp} K_{Tp}^T,
$$

where $K_{Tp}$ are the temperature Jacobians calculated for the H$_2$O retrieval
range, $S_{Tp}$ includes the random errors on temperature estimated according
to the method described in subsection 4.1.2. The matrix $S_\epsilon'$ has then been
used in the retrieval of H$_2$O, and the uncertainties on the latter are calculated
from the matrix $S$.

\(^3\)The saturation vapor pressure ($P$) has been calculated by assuming an expression of
the form $P = \exp \left( a + \frac{b}{T} + cT + dT^2 \right)$, where $T$ is the temperature and the coefficients $a$, $b$, $c$,
d for H$_2$O are taken equal to 15.278, -5980.3 K, $8.8294 \times 10^{-3}$ K$^{-1}$ and $-1.2169 \times 10^{-5}$ K$^{-2}$.  

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We have then evaluated the error on the retrieved H$_2$O mole fraction coming from the uncertainty on the continuum level. For this, we have applied a shift to the continuum, representative of the uncertainty of the latter, and retrieved a new H$_2$O mole fraction. The uncertainty on the continuum has been estimated as the 1$\sigma$ significance level (68.3%) of the fit described in section 4.2, which is determined by a $\Delta \chi^2$ that depends on the number of retrieved cross-sections (Press et al., 2007).

The last source of errors propagated to the H$_2$O abundance is the temperature systematic errors coming from the choice of the CH$_4$ profile. It has been evaluated by shifting the stratospheric temperature by the systematic uncertainties estimated in section 4.1.2 following a change of $\pm 11\%$ of the CH$_4$ profile. The H$_2$O retrieval is then performed with the new temperature profile. The difference in the H$_2$O mole fraction retrieved with and without the shift of temperature is taken as the systematic uncertainty on the H$_2$O abundance. Note that the H$_2$O retrieval has also been performed using the stratospheric temperatures retrieved using a CH$_4$ mole fraction of 0.01. The impact of this change on the determined H$_2$O abundance will be discussed separately from the systematic errors.

The total uncertainty is finally calculated considering that the random uncertainties (noise, temperature and continuum) are not fully independent of each other, but are independent of the systematic ones resulting from the uncertainty on the CH$_4$ profile. The random errors have been thus linearly summed and then added to systematic errors by quadrature.

Finally, the height at which H$_2$O is retrieved is taken as the pressure level at which the Jacobians of H$_2$O are maximum, which is different for the
different assumed H$_2$O a priori profile listed in Table 3. The full width at half maximum of these functions has been used to define the uncertainty on that level.

4.3.3. Error analysis - PACS

The assessment of the error on the H$_2$O abundance retrieved from the PACS observations is very similar to the one applied for CIRS observations, apart from the evaluation of the continuum uncertainty. Because PACS spectra are scaled by a modelled continuum, the continuum level does not need to be fitted and there is no “retrieval” error on this level. However, the modelled continuum is associated with a systematic uncertainty related to the choice of the reference atmosphere, which propagates to the H$_2$O abundance. To evaluate the propagation of this uncertainty, we have to estimate the error on the modelled continuum, and thus on the parameters affecting this level. For tropospheric temperature, random errors calculated as described in section 4.1.2 are taken into account. For the aerosols distribution and the associated absorption cross-sections, based on the retrieval performed on CIRS spectra, we have estimated that their uncertainties are reasonably around 10% and 20% respectively. A new continuum level has then been modelled by scaling these parameters by their respective uncertainties, and the H$_2$O abundance has been retrieved from the new radiance spectra computed. The H$_2$O difference obtained between the two modelled continuum is taken as the systematic uncertainty on H$_2$O due to the continuum.
5. Results

5.1. Temperature

5.1.1. CIRS disc-averaged

Figure 6 (a-d) shows the a priori and retrieved temperature profiles for the four CIRS disc-averaged observations. For all of them, the retrieved profile departs from the a priori by about 1.5 K in the troposphere (at pressure levels larger than 100 mbar) and by up to 4 K in the stratosphere (between 100 and 1 mbar) for the 182, 202, 206 observations, and by up to 7 K for the set 219. Figure 5 shows the fit to the set 182 as an example.

The vertical sensitivity of the temperature retrieval can be discussed using the associated averaging kernel (AVK) functions (Figure 6 (e-h)). The AVK functions represent the sensitivity of the retrieved state \( x_{\text{ret}} \) to the true state \( x_{\text{tr}} \) (Rodgers, 2000). The height at which they reach their largest value provides an estimation of the height of the retrieval maximum sensitivity. The AVK are included in a matrix \( (A) \) calculated according to:

\[
A = \frac{\partial x_{\text{ret}}}{\partial x_{\text{tr}}} = (K^T S_e^{-1} K + S_a^{-1})^{-1} K^T S_e^{-1} K, \quad (3)
\]

where \( K \) is the Jacobian matrix of the retrieval. From Figure 6 (e-h), we can see that for the four disc-averaged observations the maximum sensitivity to temperature is found between 400 and 200 mbar as it was expected from the Jacobians (Figure 3.a). The temperature sensitivity stays large up to 70 mbar and then starts to decrease as moving upwards in the atmosphere. In the pressure range of sensitivity to \( \text{H}_2\text{O} \), the AVK functions still present non zero values and show that the retrieval is also partly sensitive to temperatures
in this region. At pressure levels lower than 1 mbar, where CH₄ lines are not
sensitive to temperature (from Jacobians and AVK), the retrieved profile
stays close to the *a priori* profile.

The random retrieval error on the temperature profile is shown in Fig-
ure 7.a, and has been calculated for the set 182. Given the very similar noise
levels and retrieval diagnostics (AVK, DOFS,...), we can reasonably assume
that the error will be of the same order of magnitude for the four sets. It
decreases from 1 K at the surface to about 0.3 K around 100 mbar, where
the maximal sensitivity to temperature is found. It increases then to 0.4 K
and stays constant up to 7 mbar, where it starts to increase again to 1 K at
1 mbar. From this level, as the sensitivity of CH₄ lines to temperature de-
creases, it increases and reaches more than 2 K. The systematic errors on the
retrieved temperature profile are shown in Figure 7.b in absolute value for
the set 182. We can see that the choice of CH₄ stratospheric abundance has a
significant impact on the retrieved temperatures. This is especially the case
in the range of sensitivity to H₂O, where the differences reach around 2 K
when the CH₄ profile is shifted by ±11% (mole fraction of CH₄ equals 1.32%
and 1.64%). In the case where the CH₄ mole fraction is reduced to 0.01, these
differences are even larger and reach more than 5 K in the 12-1 mbar range.
Note that in this case the retrieved temperature profile is warmer than the
*a priori* profile. The impact of the assumed CH₄ profile is of the same or-
der of magnitude for the three other sets. The impact of these temperature
differences on the H₂O retrieval will be assessed in subsection 5.2.1.
5.1.2. CIRS nadir

Figure 8.a shows the temperature profile retrieved from the CIRS NA spectrum and the fit is shown in Figure 5. In the troposphere, it is very similar to the *a priori* profile and this was expected since tropospheric temperatures were retrieved by Cottini et al. (2012) using CH$_4$ pure rotational lines. However, in the 12-1 mbar region, the retrieved profile is significantly different from the profile used by Cottini et al. (2012), which was taken from Anderson and Samuelson (2011) and was not retrieved from the analyzed spectrum. The differences are up to 8 K. Such temperature difference very likely impacts the retrieval of the stratospheric H$_2$O abundance and this will be investigated in subsection 5.2.2.

As seen from the AVK functions shown in Figure 8.b, the vertical sensitivity of the retrieval is very similar to the one observed for disc-averaged CIRS spectra. The AVK functions reach their maximum values between 400 and 200 mbar, where the sensitivity to temperature is maximal. Their values then start to decrease, but the retrieval still show sensitivity to temperature in the pressure range of maximum sensitivity to H$_2$O.

The random uncertainties on the temperature profile are shown in Figure 8.c. They are well below 1 K in the troposphere, especially in the range 400-30 mbar, where the sensitivity to temperature is maximal. In the region of sensitivity to H$_2$O, these errors are within 0.9-1.6 K, which is slightly larger than the ones estimated for disc-averaged observations. As already noticed in the previous subsection, the CH$_4$ profile has a large influence on the retrieved profile in the range 12-1 mbar (Figure 8.d). The systematic errors are estimated at around 2 K in absolute value if we consider ±11%
of uncertainty on the CH\textsubscript{4} profile. If the stratospheric CH\textsubscript{4} mole fraction is reduced to 0.01, the retrieved temperature profile is warmer by up to 8 K than the one shown in Figure 8.a.

5.1.3. PACS

Figure 9 presents the temperature profile retrieved from PACS observations along with the associated error and vertical sensitivity diagnostics. As observed for the CIRS NA case, the retrieved profile is significantly colder, by up to 9 K, from the temperature profile considered by Moreno et al. (2012) in the range of H\textsubscript{2}O sensitivity. As a reminder, Moreno et al. (2012) did not determine the temperature profile from PACS spectra but built one based on several retrievals performed in different studies. The observed difference could possibly lead to the retrieval of a larger H\textsubscript{2}O abundance, and then therefore to a better agreement between CIRS and Herschel measurements. This will be discussed in more details in subsection 5.2.3. The fit to the spectrum is shown in Figure 5.

From the AVK functions shown in Figure 9.b, we can see that the retrieval has less vertical resolution than those performed on CIRS observations. The maximum sensitivity to temperature is found to be within 200 and 70 mbar and thus show lower sensitivity to temperature in the lower troposphere. This was expected from the temperature Jacobians (Figure 3.c) and explains that the retrieved profile stays close to the \textit{a priori} in the low atmosphere. As observed for CIRS observations, the retrieval, however, show sensitivity to temperature in the pressure range of maximum sensitivity to H\textsubscript{2}O. The smaller vertical resolution likely explains the larger random errors on the profile in the 12-1 mbar range, which reach more than 4 K (Figure 9.c).
Systematic errors (absolute value) are within 2-3 K in the same range when the uncertainty on CH$_4$ profile is assumed to be ±11%. If the stratospheric CH$_4$ mole fraction is reduced to 0.01, the retrieved temperature profile is warmer by up to 9 K. The propagation of these uncertainties to the H$_2$O abundance will be discussed in section 5.2.3.

5.2. H$_2$O abundance

Table 4 shows a summary of the H$_2$O retrievals performed on the different observations analysed in this work. For comparison, the H$_2$O abundances retrieved when the stratospheric temperatures are determined assuming 1.00% of stratospheric CH$_4$ are also indicated. Note that they are calculated for the pressure range indicated. Example fits are shown in Figure 5. The following subsections discuss first of all the retrieval for each type of observation and then compare the retrieved H$_2$O abundances.

5.2.1. CIRS disc-averaged

5.2.1.1 CIRS 182

From Table 4, considering all the different a priori H$_2$O profiles, we can see that the H$_2$O VMR retrieved from the set 182 is found to be within 0.120 ppb and 0.215 ppb in the 8.2–5.1 mbar pressure range. This shows that the choice of the a priori profile has a significant impact on the retrieved H$_2$O abundance. The total column of H$_2$O above the condensation level is, however, less affected by the H$_2$O vertical profile assumptions and is determined to be within 3.5–3.8×10$^{14}$ molecules/cm$^2$. Because the a priori profiles have different slopes (Figure 4), they have also an effect on the height of maximum sensitivity to H$_2$O, which varies from 8.2 mbar to 5.1 mbar. For all
retrievals, the fits are reasonably good with associated $\chi^2$ around 1.48–1.49. Given these very similar values, it is therefore impossible to determine which a priori profile is the most appropriate, and all of them are kept in the following discussions. The total error has been estimated at -43%/+44%. The largest contribution to the total error is the uncertainty on the continuum level, which is estimated at ±27%. Finally, when the stratospheric temperatures retrieved using a CH$_4$ mole fraction of 0.01 are considered, the H$_2$O VMRs are around 40% smaller and fall within 0.072–0.131 ppb.

5.2.1.2 Other sets

Compared to the set 182, the three other disc-averaged spectra do not show clear H$_2$O lines as can be seen in Figure 1. Before performing the retrieval, we have first of all assessed if the presence of H$_2$O lines is statistically attested. To do this, we have calculated the misfit $\chi^2$ as a function of the scaling factor (MF) applied to the H$_2$O profile (the “IM2” a priori profile is considered), starting at MF=0. If H$_2$O emission lines are present in the spectrum, the $\chi^2$ function should show a minimum at the best fitting scaling factor. Because we are adjusting one parameter (the scaling factor), a 3σ detection requires a $\chi^2$ decrease of 9 compared to the case assuming no H$_2$O. Figure 10 shows the $\Delta\chi^2\left(\chi^2 - \chi^2(MF = 0)\right)$ functions calculated for each sets. For all of them, the minimum is found below the limit $\Delta\chi^2 = -9$, meaning that H$_2$O is detected at the 3σ level. The minimum $\chi^2$ is found at MF=0.75 for the set 202, and at MF=0.55 for the sets 206 and 219.

As the presence of H$_2$O emission lines has been statistically confirmed, we have applied the retrieval scheme to the three sets in a second stage. The
results are summarized in Table 4. The retrieved scaling factors are very close to those determined from the $\chi^2$ analysis: they are equal to 0.76, 0.58 and 0.59 respectively for the sets 202, 206 and 219. These correspond to respective H$_2$O VMRs of 0.129 ppb, 0.099 ppb, and 0.100 ppb at 5.6 mbar. As seen for the set 182, the choice of the $a$ priori H$_2$O profile has a strong influence, leading to retrieved VMRs varying within 0.080–0.151 ppb for the set 202, 0.059–0.119 ppb for the set 206, and 0.054–0.120 ppb for the set 219, mainly between 8 mbar and 5 mbar. The total columns of H$_2$O are more stable against the choice of the $a$ priori H$_2$O profile and are found to be within 2.3–2.6$\times$10$^{14}$ molecules/cm$^2$, 1.7–2.0$\times$10$^{14}$ molecules/cm$^2$, 1.6–2.1$\times$10$^{14}$ molecules/cm$^2$ respectively for sets 202, 206, and 219. The assessment of the most appropriate $a$ priori is impossible for the three sets because all profiles lead to very similar quality of fit. The reduced $\chi^2$ values are found to be 1.65–1.66, 1.40 and 1.20–1.21 respectively for sets 202, 206 and 219. The estimated total errors are larger than those assessed for the set 182. This was expected given the smaller SNR of the averaged spectra for the three other sets. These errors are equal to -56%/+57%, ±65%, and -70%/+78% respectively for sets 202, 206 and 219. Finally, as seen for set 182, when using the stratospheric temperature profile determined for a stratospheric CH$_4$ mole fraction of 0.01, the retrieved H$_2$O VMRs are smaller. They fall within 0.038–0.081 ppb, 0.038–0.082 ppb and 0.032–0.075 ppb respectively for sets 202, 206 and 219.

5.2.2. CIRS nadir

Using the scheme developed for disc-averaged CIRS spectra, the H$_2$O VMR retrieved from the CIRS NA spectrum falls within 0.058 ppb–0.109 ppb
in the 9.6–5.4 mbar pressure range. This VMR is smaller by a factor 1.3–
2.4 than the one determined by Cottini et al. (2012) (0.14 ppb) and this
shows the influence of the chosen retrieval scheme on the evaluation of the
H₂O abundance. This difference is mostly associated with the choice of the
a priori profile. Indeed, by redoing the retrieval using a constant H₂O a
priori profile of 0.1 ppb above the saturation level as Cottini et al. (2012),
we retrieve a H₂O VMR of 0.138 ppb, which is in very good agreement with
the value reported by Cottini et al. (2012). However, because the H₂O mole
fraction has been firmly demonstrated to increase with height, a H₂O profile
constant at all altitudes is not the best option for the a priori profile and
was therefore not considered in this work. Note that if the total errors are
taken into account (-42%/+44% for this work, ±0.05 ppb for Cottini et al.
(2012)), the largest H₂O abundances determined in this work agree within
the error bars with the results of Cottini et al. (2012). Note finally that the
discrepancies mentioned above are even larger when the H₂O abundance is
retrieved using the stratospheric temperatures determined for a stratospheric
CH₄ mole fraction of 0.01. The retrieved H₂O VMRs are in this case a factor
of 1.7–3.2 smaller than the one reported by Cottini et al. (2012).

As we already mentioned, the assumed temperature profile affects the
height of the H₂O condensation level and therefore the vertical sensitivity
to H₂O. Because the temperature profile we retrieved is colder than the one
assumed by Cottini et al. (2012), the maximum sensitivity to H₂O is found
to be higher in the atmosphere (9.6–5.4 mbar) than determined by Cottini
et al. (2012) (maximum sensitivity at 10.7 mbar). Note also that, as seen for
CIRS disc-averaged observations, determining the most appropriate H₂O a
priori profile is impossible based on the $\chi^2$ values (0.67–0.68 for the reduced $\chi^2$).

5.2.3. PACS

Before applying the retrieval scheme described in section 4 to the PACS observations, we have determined if differences in retrieval codes might have a significant impact on the H$_2$O abundance. For this, we first retrieved H$_2$O using very similar retrieval assumptions as those used by Moreno et al. (2012), i.e., using the same temperature profile and considering the “Sa” profile as a priori (Table 3 and Figure 4). With these settings, we retrieve a scaling factor of 0.96 to be applied to the “Sa” profile, which is in excellent agreement with the results of Moreno et al. (2012). To investigate the effect of the choice of the a priori profile, we then performed the retrievals using the different profiles shown in Table 3. The retrieved H$_2$O abundances vary between 0.015–0.031 ppb, and fall well within the error bars estimated by Moreno et al. (2012).

Since differences in retrieval codes do not affect the retrieved H$_2$O abundance, we have investigated in a second stage if differences in the retrieval method, especially differences in the temperature profile, significantly affect the retrieved H$_2$O abundance from PACS. For this, we have applied the analysis described in section 4 and the results are shown in Table 4. The retrieved VMRs are found to be within 0.024–0.052 ppb and are slightly larger than those determined by Moreno et al. (2012). This shows the impact of the temperature profile retrieved from PACS CH$_4$ lines, which was found to be colder than the one assumed by Moreno et al. (2012). This temperature profile also explains that the H$_2$O maximum sensitivity is found higher in the
atmosphere (8.2–5.6 mbar) because the condensation level moves upwards. The H$_2$O VMRs retrieved when considering the stratospheric temperature profile determined using 1.00% for the stratospheric CH$_4$ mole fraction are, however, in good agreement with those determined by Moreno et al. (2012). This is explained by the fact that this temperature profile is very similar to the one considered by Moreno et al. (2012). In this case, the retrieved H$_2$O VMRs fall within 0.020–0.030 ppb.

As already observed for the other types of observations, the H$_2$O total column is more stable against the choice of the a priori H$_2$O profile and is determined to be within 0.8–0.9×10$^{14}$ molecules/cm$^2$. The total errors are estimated to be -15%/+12%. These errors are significantly smaller than those estimated for CIRS. This can be explained by the better SNR of PACS observations. Note that the reduced $\chi^2$ values are very similar for all the a priori profiles and are within 1.49–1.51.

5.2.4. Comparison and discussion

Figure 11 compares the H$_2$O VMRs retrieved for all observations using the same analysis scheme as a function of the H$_2$O a priori profile considered. We can see that the comparison changes slightly depending on the chosen a priori, which directly impacts the retrieved abundance. The lowest VMRs are retrieved using the “Sa” profile of Moreno et al. (2012) (0.029–0.120 ppb), and the largest using the “L3” profile of Lara et al. (2014) (0.052–0.215 ppb). Because it affects the H$_2$O retrieved abundance, the choice of the a priori profile seems to impact the relative agreement between the different observations. Indeed, the best agreement is found for the “Sa” profile, whereas the retrieved VMRs show larger discrepancies in case of “L1”, “L2”, and
“L3” profiles. Note that in the case of the H$_2$O total column abundance, there is much less impact of the \textit{a priori} profile on the relative agreement between the different datasets. However, despite the differences due to the choice of the H$_2$O \textit{a priori} profile and whether the total column abundances or the volume mixing ratios are considered, the same general trend is observed. For each case, the smallest H$_2$O abundance is retrieved from PACS spectra confirming the initial discrepancy between Moreno et al. (2012) and Cottini et al. (2012). The CIRS disc-averaged set 182 is associated with the largest H$_2$O abundance, and a very good agreement is found between the three other disc-averaged sets and the CIRS NA observations. Note that the \textit{a priori} profile also impacts the height of maximum sensitivity (Table 4).

By analysing PACS observations with the same assumptions than Moreno et al. (2012) but with the NEMESIS retrieval code, we have shown that differences in the retrieval codes used by Moreno et al. (2012) and Cottini et al. (2012) hardly explain the differences observed previously between CIRS and PACS measurements. Applying the same retrieval method and making the same H$_2$O \textit{a priori} assumptions for both PACS and CIRS observations has, on the contrary, decreased the discrepancies observed previously between these instruments (Moreno et al., 2012; Cottini et al., 2012). However, the analysis of the CIRS disc-averaged dataset has brought new differences. Whereas the sets 202, 206, and 219 are in good agreement with the CIRS
NA observation and consistent within the error bars with PACS measurements, the H$_2$O VMRs retrieved from the set 182 are significantly larger and especially differ from PACS measurements. This is confirmed by Figure 12, which shows a comparison between observed PACS H$_2$O lines and lines simulated using the H$_2$O profiles retrieved for the different observations considering the “Sa” a priori profile. Indeed, we can see that the profile retrieved for set 182 strongly overestimates PACS H$_2$O lines. Also from Figure 12 we can see that, although a relative agreement is found within the error bars between the other CIRS and PACS measurements, the H$_2$O profiles retrieved from CIRS observations are inconsistent with PACS H$_2$O lines. This confirms that H$_2$O is less abundant for PACS observations. Based on this, the hypothesis of natural variations of the H$_2$O abundance, triggered by the global circulation in Titan’s atmosphere, can be envisaged. Similarly to nitriles and some hydrocarbons, H$_2$O shows an increasing mole fraction with height, resulting from a source at high altitudes (either micrometeorite ablation or Enceladus) and a sink in the low atmosphere (condensation and photolysis). The meridional circulation could thus possibly affect the latitudinal distribution of H$_2$O abundance. This has been observed for nitriles and hydrocarbons, for which subsidence of air rich in these gases leads to an enrichment in the stratosphere, and inversely upwelling leads to the depletion of their stratospheric abundance (e.g. Coustenis et al., 2007, 2016; Teanby et al., 2006, 2008, 2009, 2010). The observations analysed in this work are recorded at different sub-spacecraft latitudes and at different periods in Titan’s year (see Figure 11), and are therefore subject to different meridional circulation. The CIRS disc-averaged set 182 is recorded in the
late southern autumn-early southern winter, when according to dynamical
models (e.g. Hourdin et al., 2004) the meridional circulation is dominated
by a single Hadley cell with subsidence at the South Pole (Teanby et al.,
2012) and upwelling at the North Pole. The sub-spacecraft latitude of this
set is -41° but because of its disc-averaged nature, the observation includes
the South Pole. It is therefore sensitive to the possible H$_2$O enhancement
brought by the subsidence of air rich in H$_2$O occuring at the South Pole,
and this could explain the larger stratospheric VMRs retrieved for this set
(see also Coustenis et al. (2016)). On the contrary, PACS observations were
recorded close to the northern spring equinox, for which dynamical models
predict a transitional circulation state, in which two cells develop with up-
wellings at the equator and subsidence at both poles (Hourdin et al., 2004;
Vinatier et al., 2015). The PACS sub-spacecraft latitude is close to the equa-
tor, where upwelling could cause H$_2$O depletion and explain the low H$_2$O
abundance retrieved from PACS spectra. Note that the variations of H$_2$O
could be very similar to those observed for C$_2$H$_2$, which has a comparable
lifetime of around 10 years. Another origin that could contribute to explain
the H$_2$O VMRs differences is the spatial/time variation of the input OH/H$_2$O
flux in Titan’s atmosphere. Depending on its source, H$_2$O is likely deposited
at different locations (heights but also latitudes) in Titan’s atmosphere. The
idea of time variations on scales of decades/centuries has already been put
forward by Moreno et al. (2012) to explain the difference of the input flux
needed to explain current H$_2$O and CO$_2$ abundances. Time variations of
the H$_2$O source rate at Enceladus, although previously debated, have been
reported by some recent studies (Hedman et al., 2013; Nimmo et al., 2014;
Teolis et al., 2017, e.g.). These variations occur on timescales of a few months or years, which is smaller than the H$_2$O lifetime, and unlikely explain the H$_2$O differences observed in this work. Variations on longer timescales cannot be ruled out, but do not provide a complete satisfactory explanation for the differences observed in this work. Indeed, if such variations existed, the H$_2$O abundance would react over periods of around 10 years. The variations of the H$_2$O abundance reported in this work occur, however, on shorter timescales. Moreover, assuming that a change of the OH/H$_2$O flux is responsible for the difference between the H$_2$O mole fractions retrieved from PACS and the set 182 of CIRS, the good agreement observed between the CIRS NA measurement and the three other CIRS disc-averaged sets becomes then more difficult to understand. For this reason, the meridional circulation and its seasonal variations seem at this stage to provide the most reasonable explanation for the H$_2$O abundance differences observed in this work. Note also that longitudinal variations of H$_2$O are barely possible given the lifetime of H$_2$O and the strong stratospheric zonal winds that rapidly homogenize gas abundances. Finally, following the work of Coustenis et al. (1998), Moreno et al. (2012) reanalysed SWS/ISO observations and determined a H$_2$O VMR of 0.06 ppb. Most of the H$_2$O VMRs determined in this work are in good agreement with this reanalysis.

The hypotheses discussed above are hardly proved with the results of this work, given the very few analysed measurements and the large errors associated with the CIRS observations. The analysis of more observations will be needed in future work to better understand the observed H$_2$O variations. We also believe that the development of 2D-3D photochemical models, cou-
pling Titan’s oxygen chemistry and general circulation, could help address the new issues highlighted in this work. Note that if H$_2$O is retrieved using stratospheric temperatures determined for a stratospheric CH$_4$ mole fraction of 0.01, the H$_2$O VMRs evaluated for the set 182 and for PACS observations remain significantly different, and thus the above discussion still holds. It is worth, however, assessing if this discussion still holds when the CH$_4$ latitudinal variability is taken into account. Indeed, we have considered a constant stratospheric CH$_4$ mole fraction for all the observations (either 1.48% or 1.00%). Based on Lellouch et al. (2014), 1.48% seems to be a relatively good assumption for the CIRS disc-averaged set 182. On the contrary, a mole fraction of 1.00% is best suited for the set 219, the CIRS NA observation and for PACS data. If these CH$_4$ mole fractions are assumed for these observations, the difference in the retrieved H$_2$O VMRs between set 182 and the three other observations is larger than when considering constant CH$_4$ abundance, and this strengthens the idea of latitudinal variations of the H$_2$O abundance. Regarding the sets 202 and 206, it is more difficult to assess what CH$_4$ mole fraction is the most appropriate, given the large error bars reported by Lellouch et al. (2014). If it is assumed to be in between the extreme values considered in this work, significant differences in the retrieved H$_2$O VMRs should remain. In future work, attempts to reduce the uncertainty on the stratospheric CH$_4$ mole fraction, but also on other parameters such as the continuum level, should be undertaken to reduce the total uncertainty on the retrieved H$_2$O abundance.
6. Conclusions

The goal of this work was to understand the differences in H$_2$O abundances retrieved previously by Moreno et al. (2012) from Herschel disc-averaged observations, and by Cottini et al. (2012) from the average of a large number of spectra recorded in a 0–30° latitudinal bin. The strategy for this was to 1) analyse disc-averaged observations of CIRS to investigate if the observation geometry could explain the differences, and 2) (re)analyse the three types of observation with the same retrieval scheme to assess if differences in retrieval codes/methodology could be responsible for the previous discrepancies. For this purpose, four sets of CIRS disc-averaged observations taken between 2013–2015 were analysed. A retrieval scheme composed of three successive steps was also set up. It comprises the retrieval of the temperature profile (step 1), the adjustment of the continuum level in the vicinity of H$_2$O lines (step 2), and the retrieval of the H$_2$O abundance (step 3). A rigorous error assessment was performed and included the propagation of the uncertainties of key parameters: measurement noise, temperature, CH$_4$ profile (for temperature retrieval), and continuum level. Different H$_2$O a priori profiles, either consistent with PACS or CIRS measurements, have been considered for the H$_2$O retrieval from all observations.

With this analysis, we have first of all shown that differences in the retrieval codes used by Moreno et al. (2012) and Cottini et al. (2012) do not explain the discrepancies observed previously between PACS and CIRS measurements. Applying the same retrieval method, including the temperature retrieval, and making the same H$_2$O a priori assumptions for both PACS and CIRS has, on the contrary, better reconciled the measurements of these in-
struments, with the agreement depending on the chosen \textit{a priori} H$_2$O profile. The retrieved H$_2$O abundances are also in good agreement with the reanalysis of ISO observations performed by Moreno et al. (2012). The addition of the disc-averaged CIRS observations, instead of confirming the consistency between the different datasets, has revealed discrepancies mainly between the set 182 and PACS observations. This makes us question the idea that the datasets should be in agreement. The instruments measured H$_2$O at different sub-observer latitudes and at different periods in Titan’s year which are associated with different meridional circulation. This circulation is shown to affect the latitudinal distribution of nitriles and hydrocarbons, through the subsidence or upwelling of air rich in these gases. As discussed in this work, the meridional circulation could potentially impact the latitudinal distribution of H$_2$O as well and explain the differences between the CIRS set 182 and PACS measurements. On the contrary, time variations of the OH/H$_2$O input flux in Titan’s atmosphere, which were shown to occur on timescales of a few months or years by recent studies, unlikely explain the observed differences given the H$_2$O lifetime of around 10 years. Given the few H$_2$O measurements and the large uncertainties associated with the CIRS observations, these hypotheses are, however, impossible to evaluate with the results obtained so far.

To validate these assumptions, the analysis of more observations will be required in future work. More especially, the NASA Cassini mission covers now almost the half of a Titan year. CIRS FP1 observations of Titan recorded between 2004 and 2017 could be used to determine if the H$_2$O abundance varies with time and/or latitude. The photochemical models describing the
oxygen chemistry in Titan’s atmosphere developed until now are 1D models, considering only the vertical transport of gases. In the future 2D-3D models, coupling the oxygen chemistry with the general circulation, should be developed to address the questions arising from this work. Finally, efforts should be undertaken to decrease the uncertainties on the retrieved H$_2$O abundances, which are revealed to be large for CIRS observations. Reducing the uncertainty of the stratospheric CH$_4$ mole fraction or on the continuum level, which were shown to significantly affects the H$_2$O retrieval, should be a priority of any future study.

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Table 1: Sets of CIRS FP1 disc-averaged observations analysed in this work. The start and end times are in UTC and Titan radius is the projected angular radius of the solid surface. The sub-spacecraft latitudes and Titan radius are the average of each sequence. TEA stands for Titan Explorations at Apoapse.

<table>
<thead>
<tr>
<th>TEA</th>
<th>Start date</th>
<th>End date</th>
<th>No. spectra</th>
<th>Sub-spacecraft latitude</th>
<th>Titan radius</th>
</tr>
</thead>
<tbody>
<tr>
<td>182-001</td>
<td>21/02/2013</td>
<td>22/02/2013</td>
<td>712</td>
<td>-41.43°</td>
<td>1.64</td>
</tr>
<tr>
<td></td>
<td>08:31:38</td>
<td>04:21:23</td>
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</tr>
<tr>
<td>202-001</td>
<td>02/03/2014</td>
<td>03/03/2014</td>
<td>898</td>
<td>50.73°</td>
<td>1.69</td>
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<tr>
<td></td>
<td>22:47:21</td>
<td>11:56:10</td>
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<td></td>
<td></td>
</tr>
<tr>
<td>206-001</td>
<td>10/07/2014</td>
<td>10/07/2014</td>
<td>793</td>
<td>48.58°</td>
<td>1.67</td>
</tr>
<tr>
<td></td>
<td>00:50:30</td>
<td>12:29:06</td>
<td></td>
<td></td>
<td></td>
</tr>
<tr>
<td>219-001</td>
<td>23/07/2015</td>
<td>24/07/2015</td>
<td>792</td>
<td>-0.02°</td>
<td>1.63</td>
</tr>
<tr>
<td></td>
<td>13:57:21</td>
<td>01:35:55</td>
<td></td>
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</tbody>
</table>
Table 2: PACS observations of CH₄ and H₂O (see also Moreno et al. (2012)) analysed in this work.

<table>
<thead>
<tr>
<th>Gas</th>
<th>Line (µm)</th>
<th>Spectral resolution (µm)</th>
<th>Start date</th>
<th>Duration (sec)</th>
<th>Sub-observer lat/lon</th>
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</thead>
<tbody>
<tr>
<td>H₂O</td>
<td>66.43</td>
<td>0.0169</td>
<td>22/06/2010</td>
<td>3071</td>
<td>2.10° / 283.74°</td>
</tr>
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<td></td>
<td></td>
<td></td>
<td></td>
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<td>23:49:52</td>
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<tr>
<td></td>
<td>75.38</td>
<td>0.039</td>
<td>22/06/2010</td>
<td>3079</td>
<td>2.10° / 282.91°</td>
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<tr>
<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td>22:56:26</td>
</tr>
<tr>
<td></td>
<td>108.07</td>
<td>0.1127</td>
<td>22/06/2010</td>
<td>3079</td>
<td>2.10° / 282.91°</td>
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<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td>22:56:26</td>
</tr>
<tr>
<td>CH₄</td>
<td>106.43</td>
<td>0.1120</td>
<td>23/06/2010</td>
<td>1140</td>
<td>2.10° / 284.59°</td>
</tr>
<tr>
<td></td>
<td></td>
<td></td>
<td></td>
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<td>00:43:14</td>
</tr>
<tr>
<td></td>
<td>119.63</td>
<td>0.1176</td>
<td>23/06/2010</td>
<td>1140</td>
<td>2.10° / 284.59°</td>
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<td></td>
<td></td>
<td></td>
<td></td>
<td></td>
<td>00:43:14</td>
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</table>
Table 3: *A priori* profiles considered for the H$_2$O retrieval (see also Figure 4).

<table>
<thead>
<tr>
<th>Reference</th>
<th>Name</th>
<th>Type</th>
<th>Agreement</th>
<th>Comments</th>
</tr>
</thead>
<tbody>
<tr>
<td>Moreno et al. (2012)</td>
<td>Sa</td>
<td>Measured</td>
<td>Herschel</td>
<td>Parametrized profile: ( q = q_0 \times (p_0/p)^n ) with ( n ) the power law index</td>
</tr>
<tr>
<td>Dobrijevic et al. (2014)</td>
<td>IM1</td>
<td>Modelled</td>
<td>Herschel</td>
<td>Origin H$_2$O: micrometeorite ablation OH flux: ( 5.7 \times 10^7 \text{ cm}^{-2} \text{ s}^{-1} )</td>
</tr>
<tr>
<td>Dobrijevic et al. (2014)</td>
<td>IM2</td>
<td>Modelled</td>
<td>CIRS</td>
<td>Origin H$_2$O: micrometeorite ablation OH flux: ( 2.6 \times 10^6 \text{ cm}^{-2} \text{ s}^{-1} )</td>
</tr>
<tr>
<td>Dobrijevic et al. (2014)</td>
<td>IE1</td>
<td>Modelled</td>
<td>Herschel</td>
<td>Origin H$_2$O: Enceladus OH flux: ( 1.4 \times 10^6 \text{ cm}^{-2} \text{ s}^{-1} )</td>
</tr>
<tr>
<td>Lara et al. (2014)</td>
<td>L1</td>
<td>Modelled</td>
<td>Herschel</td>
<td>Time-dependent model, initial OH flux: ( 5.1 \times 10^6 \text{ cm}^{-2} \text{ s}^{-1} + \text{loss to the haze} )</td>
</tr>
<tr>
<td>Lara et al. (2014)</td>
<td>L2</td>
<td>Modelled</td>
<td>CIRS</td>
<td>Time-dependent model, initial OH flux: ( 5.1 \times 10^6 \text{ cm}^{-2} \text{ s}^{-1} + \text{loss to the haze} )</td>
</tr>
<tr>
<td>Lara et al. (2014)</td>
<td>L3</td>
<td>Modelled</td>
<td>CIRS</td>
<td>Time-dependent model, initial OH flux: ( 1.0 \text{ or } 2.0 \times 10^7 \text{ cm}^{-2} \text{ s}^{-1} )</td>
</tr>
</tbody>
</table>

Table 4: Retrieved H$_2$O abundances and associated errors for the different observations analysed. The retrieved abundances are indicated for the two temperature profiles considered in this work. The given ranges describe the results arising from using the different *a priori* H$_2$O profiles listed in Table 3. Note that the pressure range is defined using the pressure levels at which the H$_2$O sensitivity is maximum (based on H$_2$O Jacobians). The total column of H$_2$O is calculated above the condensation level. For the comparison, the reported H$_2$O VMRs for the case using 1.00% of stratospheric CH$_4$ are given for the same pressure range.

<table>
<thead>
<tr>
<th>Observation</th>
<th>H$_2$O VMR (ppb)</th>
<th>Total error</th>
<th>Pressure (mbar)</th>
<th>H$_2$O total column (mole/cm$^2$) if 1.00% CH$_4$</th>
<th>H$_2$O VMR (ppb)</th>
<th>Previous retrieved H$_2$O</th>
</tr>
</thead>
<tbody>
<tr>
<td>CIRS 182</td>
<td>0.120-0.215</td>
<td>-43%/+44%</td>
<td>8.2-5.1</td>
<td>3.5-3.8×10$^{14}$</td>
<td>0.072-0.131</td>
<td>/</td>
</tr>
<tr>
<td>CIRS 202</td>
<td>0.080-0.151</td>
<td>-56%/+57%</td>
<td>8.2-5.1</td>
<td>2.3-2.6×10$^{14}$</td>
<td>0.038-0.081</td>
<td>/</td>
</tr>
<tr>
<td>CIRS 206</td>
<td>0.059-0.119</td>
<td>-65%/+65%</td>
<td>8.0-5.1</td>
<td>1.7-2.0×10$^{14}$</td>
<td>0.038-0.082</td>
<td>/</td>
</tr>
<tr>
<td>CIRS 219</td>
<td>0.054-0.120</td>
<td>-70%/+78%</td>
<td>8.2-4.7</td>
<td>1.6-2.1×10$^{14}$</td>
<td>0.032-0.075</td>
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<tr>
<td>CIRS NA</td>
<td>0.058-0.109</td>
<td>-42%/+44%</td>
<td>9.6-5.4</td>
<td>2.0-2.2×10$^{14}$</td>
<td>0.044-0.082</td>
<td>0.14±0.05 ppb at 10.7 mbar</td>
</tr>
<tr>
<td>PACS</td>
<td>0.024-0.052</td>
<td>-15%/+12%</td>
<td>8.2-5.6</td>
<td>0.8-0.9×10$^{14}$</td>
<td>0.020-0.030</td>
<td>0.023±0.006 ppb at 12.1 mbar</td>
</tr>
</tbody>
</table>
Figure 1: CIRS and PACS spectra analysed in this work. For display purposes, CIRS disc-averaged spectra have been offset vertically as indicated. H$_2$O lines, which are used for the retrieval, are identified for CIRS spectra with vertical dashed lines. PACS spectra are expressed in radiance units (µW cm$^{-2}$ sr$^{-1}$ µm$^{-1}$) as explained in the text.
Figure 2: Main reference atmospheric vertical profiles used to simulate CIRS and PACS spectra: (a) Mole fraction of gases (Niemann et al., 2010), (b) Haze specific density (de Kok et al., 2007b, 2010b; Tomasko et al., 2008). See text for details.
Figure 3: Jacobians of temperature and H$_2$O (color scales) calculated for disc-averaged observations of CIRS (a,b) and PACS (c,d). The Jacobians of H$_2$O are calculated for the “IM2” a priori profile (see Table 3 and Figure 4). For CIRS, the H$_2$O lines used in the retrieval are indicated with the vertical black lines.
Figure 4: *A priori* profiles of H$_2$O considered in this work (see Table 3 for details).
Figure 5: Example fits to the observed CIRS (set 182 and CIRS NA) and PACS spectra. For CIRS observations, they include the fits of the temperature retrieval, of the continuum and of the H₂O retrieval (considering the “IM2” *a priori* profile). Vertical lines indicate the fitted H₂O lines. For display purposes, the CIRS NA spectrum has been offset vertically as indicated. For PACS, the fits are shown for the H₂O and CH₄ lines analysed (considering the “Sa” *a priori* profile).
Figure 6: Retrieved and *a priori* temperature profiles (a-d) and associated averaging kernels functions (e-h) for the four sets of disc-averaged CIRS spectra. The black dashed lines delimit the region of maximum sensitivity to H$_2$O (12-1 mbar). For clarity, only one in every two averaging kernel functions until 0.2 mbar are shown. The dots indicate the pressure levels associated to each of these functions and a different colour is associated to each AVK function.
Figure 7: (a) Random error and its standard deviation calculated from retrievals on synthetic spectra. (b) Systematic errors (absolute values) on the retrieved temperature profile associated with the choice of CH$_4$ profile. The black dashed lines delimit the stratospheric region of sensitivity to H$_2$O abundance.
Figure 8: Temperature retrieval from the CIRS NA spectrum: (a) A priori and retrieved profile, (b) AVK functions associated to the retrieved profile, (c) Random errors and its standard deviation calculated from retrievals on synthetic spectra, and (d) Systematic errors (absolute values) on the retrieved temperature profile associated to the choice of CH$_4$ profile. The black dashed lines delimit the stratospheric region of sensitivity to H$_2$O abundance.
Figure 9: Same as Figure 8 but for PACS observations.
Figure 10: $\Delta \chi^2(\chi^2 - \chi^2(MF=0))$ as a function of the scaling factor applied to the “IM2” H$_2$O a priori profile for the CIRS sets 202 (a), 206 (b), and 219 (c). The horizontal dashed line indicates the level of $3\sigma$ detection ($\Delta \chi^2 = -9$).
Figure 11: Retrieved $\text{H}_2\text{O}$ VMRs for all observations as a function of latitude, of time (color scale) and of the \textit{a priori} $\text{H}_2\text{O}$ profile. These VMRs are retrieved using the stratospheric temperature profile determined using 1.48\% for the stratospheric $\text{CH}_4$ mole fraction. The circles represent CIRS disc-averaged observations, the upward-pointing triangle the CIRS nadir observation, and the left-pointing triangle PACS observations.
Figure 12: Comparison between observed PACS H$_2$O lines (black) and lines simulated using the H$_2$O profiles retrieved for the different observations (colored lines). The case using the “Sa” a priori profile has been considered (the profile has been scaled by the retrieved scaling factors) and using the stratospheric temperature profile determined using a stratospheric CH$_4$ mole fraction of 1.48%.